

# Analysing Globular Cluster observations

## Models and Analysis Tools for Lick/IDS indices

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Received xxx / Accepted xxx

**Abstract.** We have extended our evolutionary synthesis code, *GALEV*, to include Lick/IDS absorption-line indices for both simple and composite stellar population models (star clusters and galaxies), using the polynomial fitting functions of Worthey et al. (1994) and Worthey & Ottaviani (1997). We present a mathematically advanced Lick Index Analysis Tool (*LINO*) for the determination of ages and metallicities of globular clusters (GCs): An extensive grid of *GALEV* models for the evolution of star clusters at various metallicities over a Hubble time is compared to observed sets of Lick indices of varying completeness and precision. A dedicated  $\chi^2$  - minimisation procedure selects the best model including  $1\sigma$  uncertainties on age and metallicity. We discuss the age and metallicity sensitivities of individual indices and show that these sensitivities themselves depend on age and metallicity; thus, we extend Worthey's (1994) concept of a "metallicity sensitivity parameter" for an *old* stellar population at *solar* metallicity to younger clusters of different metallicities. We find that indices at low metallicity are generally more age sensitive than at high metallicity. Our aim is to provide a robust and reliable tool for the interpretation of star cluster spectra becoming available from 10m class telescopes in a large variety of galaxies – metal-rich & metal-poor, starburst, post-burst, and dynamically young. We test our analysis tool using observations from various authors for Galactic and M31 GCs, for which reliable age and metallicity determinations are available in the literature, and discuss in how far the observational availability of various subsets of Lick indices affects the results. For M31 GCs, we discuss the influence of non-solar abundance ratios on our results.

All models are accessible from our website, <http://www.astro.physik.uni-goettingen.de/~galev/>.

**Key words.** globular clusters: general – methods: data analysis – techniques: spectroscopic – galaxies: star clusters – galaxies: individual (Galaxy, M31)

### 1. Introduction

In order to understand the formation and evolution of galaxies, one of the essential issues is to reveal their star formation histories (SFHs). Unfortunately, most galaxies are observable in integrated light only, so that SFH determinations using the most reliable CMD approach are possible only for a very limited sample of nearby galaxies. However, the age and metallicity distributions of star cluster and globular cluster (GC) systems can provide important clues about the formation and evolutionary history of their parent galaxies. E.g., the violent formation history of elliptical galaxies as predicted from hierarchical or merger scenarios, is, in fact, more directly obtained from the age and metallicity distributions of their GC systems than from their integrated spectra that are always dominated by stars originating from the last major star formation episode. By means of evolutionary synthesis models, for example, we can

show that, using the integrated light of a galaxy's (composite) stellar content alone, it is impossible to date (and actually to identify) even a very strong starburst if this event took place more than two or three Gyrs ago (Lilly 2003, Lilly & Fritze – von Alvensleben 2005). Therefore, it is an important first step towards the understanding of the formation and evolution of galaxies to constrain the age and metallicity distributions of their star cluster systems (for recent reviews see, for example, Kissler-Patig 2000 and Zepf 1999, 2002), as well as of their stars (see, e.g., Harris et al. 1999, Harris & Harris 2000, 2002). Star clusters can be observed one-by-one to fairly high precision in galaxies out to Virgo cluster distances, even on bright and variable galaxy backgrounds, both in terms of multi-band imaging and in terms of intermediate resolution spectroscopy. For young star cluster systems, we have shown that the age and metallicity distributions can be obtained from a comparison of multi-band imaging data with a grid of model SEDs using the SED Analysis Tool *AnalySED* (Anders et al. 2004).

Our aim is to extend the analysis of star cluster age and metallicity distributions in terms of parent galaxy formation histories and scenarios to intermediate age and old star cluster systems. However, since for all colors the evolution slows down considerably at ages older than about 8 Gyr, even with several passbands and a long wavelength basis the results are severely uncertain for old GCs; colors – even when combining optical and near infrared – do not allow to completely disentangle the age-metallicity degeneracy (cf. Anders et al. 2004). Absorption line indices, on the other hand, are a promising tool towards independent and more precise constraints on ages and metallicities. Therefore, we present a grid of new evolutionary synthesis models for star clusters including Lick/IDS indices complementing the broad band colors and spectra of our previous models, and a Lick Index Analysis Tool *LINO* meant to complement our SED Analysis Tool. With these two analysis tools, we now possess reasonable procedures for the interpretation of both broad-band color and spectral index observations.

In an earlier study, we already incorporated a subset of Lick indices into our evolutionary synthesis code *GALEV* (Kurth et al. 1999). However, since then the input physics for the code changed considerably; instead of the older tracks we are now using up-to-date Padova isochrones, which include the thermally pulsing asymptotic giant branch (TP-AGB) phase of stellar evolution (see Schulz et al. 2002). In this work, we present the integration of the full set of Lick indices into our code. Now, our *GALEV* models consistently describe the time evolution of spectra, broad-band colors, emission lines and Lick indices for both globular clusters (treated as single-age single-metallicity, i.e. “simple” stellar populations, SSPs) and galaxies (composite stellar populations, CSPs), using the same input physics for all models (for an exhaustive description of *GALEV* and its possibilities, as well as for recent extensions of the code and its input physics, see Schulz et al. 2002, Anders & Fritze – v. Alvensleben 2003, and Bicker et al. 2004).

A recent publication (Proctor et al. 2004) also presented an analysis tool for Lick indices using a  $\chi^2$ -approach. However, they do not provide any confidence intervals for their best fitting models. In this respect, our new tool extends their approach. A drawback of our models is that, at the present stage, they do not account for variations in  $\alpha$ -enhancement, as Proctor et al. (2004) do. However, our analysis tool *LINO* is easily applicable to any available set of absorption line indices.

In Section 2, we recall the basic definitions of Lick indices, and describe how we synthesize them in our models; we also address non-solar abundance ratios. Some examples for SSP model indices are presented and briefly confronted with observations. In Section 3, Worthey’s (1994) “metallicity sensitivity parameter” is discussed and extended from old stellar populations to stellar populations of all ages. Section 4 describes and tests our new Lick Index Analysis Tool; Galactic and M31 globular cluster observations are analysed and compared with results (taken from literature) from reliable CMD analysis, and from index analyses using models with varying  $\alpha$ -enhancements, respectively. Section 5 summarizes the results and provides an outlook.

## 2. Models and input physics

### 2.1. Evolutionary synthesis of Lick indices

Lick indices are relatively broad spectral features, and robust to measure. They are named after the most prominent absorption line in the respective index’s passband. However, this does not necessarily mean that a certain index’s strength is exclusively or even dominantly due to line(s) of this element. (see, e.g., Tripicco & Bell 1995). Beyond the fact that more than one line can be present in the index’s passband, strong lines in the pseudo-continua can also affect the index strength. Most indices are given in units of their equivalent width (EW) measured in Å:

$$\text{EW}[\text{\AA}] = \int_{\lambda_1}^{\lambda_2} \frac{F_C(\lambda) - F_I(\lambda)}{F_C(\lambda)} d\lambda, \quad (1)$$

whereas index strengths of broad molecular lines are given in magnitudes:

$$I[\text{mag}] = -2.5 \log \left[ \left( \frac{1}{\lambda_1 - \lambda_2} \right) \int_{\lambda_1}^{\lambda_2} \frac{F_I(\lambda)}{F_C(\lambda)} d\lambda \right]. \quad (2)$$

$F_I(\lambda)$  is the flux in the index covering the wavelength range between  $\lambda_1$  and  $\lambda_2$ ;  $F_C(\lambda)$  is the continuum flux defined by two “pseudo-continua” flanking the central index passband.

There are currently 25 Lick indices, all within the optical wavelength range:  $\text{H}\delta_A$ ,  $\text{H}\gamma_A$ ,  $\text{H}\delta_F$ ,  $\text{H}\gamma_F$ ,  $\text{CN}_1$ ,  $\text{CN}_2$ ,  $\text{Ca}4227$ ,  $\text{G}4300$ ,  $\text{Fe}4383$ ,  $\text{Ca}4455$ ,  $\text{Fe}4531$ ,  $\text{Fe}4668$ ,  $\text{H}\beta$ ,  $\text{Fe}5015$ ,  $\text{Mg}_1$ ,  $\text{Mg}_2$ ,  $\text{Mgb}$ ,  $\text{Fe}5270$ ,  $\text{Fe}5335$ ,  $\text{Fe}4506$ ,  $\text{Fe}5709$ ,  $\text{Fe}5782$ ,  $\text{Na D}$ ,  $\text{TiO}_1$ , and  $\text{TiO}_2$ . For a full description and all index definitions, see Trager et al. (1998) and references therein.

As the basis for our evolutionary synthesis models we employ the polynomial fitting functions of Worthey et al. (1994) and Worthey & Ottaviani (1997), which give Lick index strengths of individual stars as a function of their effective temperature  $T_{\text{eff}}$ , surface gravity  $g$ , and metallicity  $[\text{Fe}/\text{H}]$ . Worthey et al. have calibrated their fitting functions empirically using solar-neighbourhood stars.

Model *uncertainties* are calculated as follows (Worthey 2004):

$$\sigma_{\text{model}} = \frac{\sigma_{\text{star}} \times \text{RMS}_{\text{fit}}}{\sqrt{N}} \quad (3)$$

with  $\sigma_{\text{star}}$  being the typical rms error per observation for the calibration stars and  $\text{RMS}_{\text{fit}}$  being the residual rms of the fitting functions (both values are given in Worthey et al. (1994) and Worthey & Ottaviani (1997), respectively).  $N$  is the number of stars in the “neighbourhood” of the fitting functions in the  $T_{\text{eff}}$ ,  $g$ ,  $[\text{Fe}/\text{H}]$  space, which is typically of the order of 25. Note that this approach is only an approximation; the real model error is most likely a strong function of  $T_{\text{eff}}$ ,  $g$ , and  $[\text{Fe}/\text{H}]$ .

Other input physics of our models include the theoretical spectral library from Lejeune et al. (1997, 1998) as well as theoretical isochrones from the Padova group for  $Z=0.0004$ , 0.004, 0.008, 0.02 and 0.05 (cf. Bertelli et al. 1994), and for  $Z=0.0001$  (cf. Girardi et al. 1996); recent versions of these isochrones include the TP-AGB phase of stellar evolution (not presented in the referenced papers) which is important for

intermediate age stellar populations (cf. Schulz et al. 2002). We assume a standard Salpeter (1955) initial mass function (IMF) from 0.15 to about 70  $M_{\odot}$ ; lowest mass stars ( $M_{\odot} < 0.6$ ) are taken from Chabrier and Baraffe (1997) (cf. Schulz et al. 2002 for details). Throughout this paper, we identify the metallicity  $Z$  with  $[\text{Fe}/\text{H}]$  and define  $[\text{Fe}/\text{H}] := \log(Z/Z_{\odot})$ .

To calculate the time evolution of Lick indices for SSP or galaxy models, we follow four steps:

1. We use the values for  $T_{\text{eff}}$ ,  $g$ , and  $[\text{Fe}/\text{H}]$  given (directly or indirectly) by the isochrones to compute the index strength  $\text{EW}_{\text{star}}$  or  $I_{\text{star}}$  for each star along the isochrones.
2. A spectrum is assigned to each star on a given isochrone and used to compute its continuum flux  $F_C$ <sup>1</sup>.
3. For each isochrone, the index strengths are integrated over all stellar masses  $m$  (after transformation of the index strengths into fluxes), weighted by the IMF (using a weighting factor  $w$ ):

$$\text{EW}_{\text{SSP}} = (\lambda_2 - \lambda_1) \cdot \left( 1 - \frac{\sum_m (F_I \cdot w)}{\sum_m (F_C \cdot w)} \right), \quad (4)$$

with  $F_I$  being a function of  $\text{EW}_{\text{star}}$  and  $F_C$ :

$$F_I = F_C \cdot \left( 1 - \frac{\text{EW}_{\text{star}}}{\lambda_2 - \lambda_1} \right). \quad (5)$$

The result is a grid of SSP models for all available isochrones, i.e., for 50 ages between 4 Myr and 20 Gyr, and the 6 metallicities given above.

4. For each time step in the computation of a stellar population model, our evolutionary synthesis code *GALEV* gives the contribution of each isochrone to the total population. To obtain galaxy model indices (or a better age resolution for SSP models), we integrate our grid of SSP models using equations (4) and (5) again, but with  $w$  being the isochrone contribution as a new weighting factor (now doing the summation over all isochrones instead of all masses),  $F_C$  being the integrated continuum flux level for each isochrone, and using  $\text{EW}_{\text{SSP}}$  instead of  $\text{EW}_{\text{star}}$ .

Following this way, we have computed a large grid of SSP models, consisting of 6 metallicities and 4000 ages from 4 Myr to 16 Gyr in steps of 4 Myr; each point of the model grid consists of all 25 Lick indices currently available.

## 2.2. Non-solar abundance ratios

Abundance ratios reflect the relation between the characteristic time scale of star formation and the time scales for the release of, e.g., SNe II products (Mg and other  $\alpha$ -elements), SNe Ia products (Fe), or nucleosynthetic products from intermediate-mass stars (N). Galaxies with different SFHs will hence be characterised by different distributions of stellar abundances ratios. This means, the Galactic relation between abundance ratios and metallicity (Edvardsson et al. 1993, Pagel &

Tautvaišienė 1995) is not necessarily valid for galaxies of different types and formation histories. Empirical index calibrations based on Galactic stars, like the fitting functions from Worthey et al. (1994) and Worthey & Ottaviani (1997) applied in this work, are based on the implicit inclusion of the Galactic relation between abundance ratios and metallicity.

A lot of work has been done in the last years to study the impact of  $\alpha$ -enhancement on stellar population models and their applications: E.g., based mainly on the work of Tripicco & Bell (1995) and Trager et al. (2000a), Thomas et al. (2003, 2004) presented SSP models of Lick indices with variable abundance ratios that are corrected for the bias mentioned above, providing for the first time well-defined  $[\alpha/\text{Fe}]$  ratios at all metallicities. The impact of these new models on age and metallicity estimates of early type galaxies is investigated in detail by Maraston et al. (2003), Thomas & Maraston (2003), Thomas et al. (2004), as well as by Trager et al. (2000a,b), among others.

However, since our purpose is to present consistently computed models for spectra, colors, emission lines *and* Lick indices for both SSPs and CSPs, a consistent attempt to allow our evolutionary synthesis code *GALEV* to account for arbitrary abundance ratios would have to be based on stellar evolutionary tracks or isochrones, detailed nucleosynthetic stellar yields, and model atmospheres for various abundance ratios. Since both consistent and complete datasets of this kind are not yet available (though first sets of evolutionary tracks for stars with enhanced  $[\alpha/\text{Fe}]$  ratios were presented by Salasnich et al. 2000 and Kim et al. 2002), at the present stage our models do *not* explicitly allow for variations in  $\alpha$ -enhancement. This is an important caveat to be kept in mind for the interpretation of extra-galactic GC populations. We think that the extensive studies on non-solar abundance ratios cited above will allow to estimate the impact of this caveat on our results.

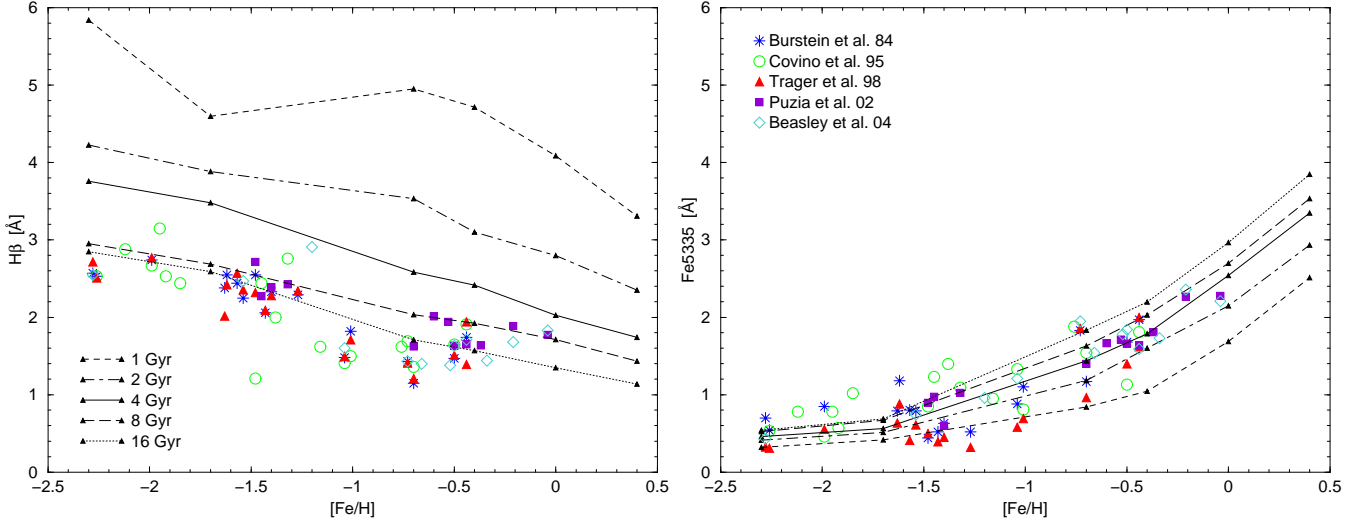
However, in Section 4.3 we show that our method is robust enough to give very good age and metallicity determinations for GCs even without using  $\alpha$ -enhanced models.

## 2.3. SSP model indices: Some examples

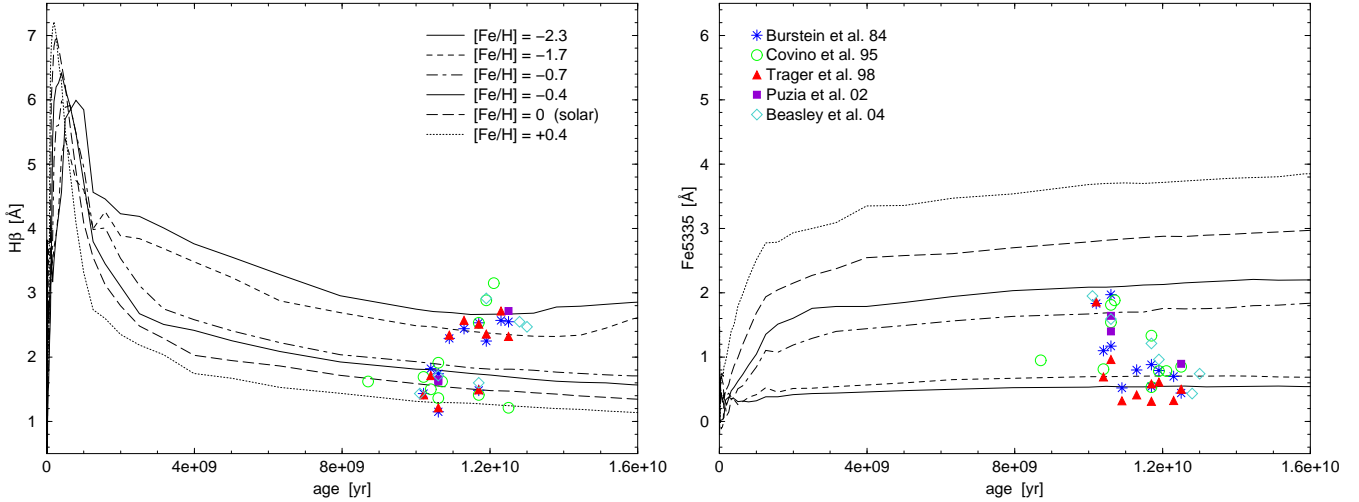
In Figs. 1 and 2, we show the time evolution and metallicity dependence of the indices  $\text{H}\beta$  and  $\text{Fe}5335$  in our new SSP models, and compare them with index measurements of Galactic GCs that are plotted against reliable age and metallicity estimates, respectively.

In particular, in Fig. 1 we confront SSP models for five ages between 1 and 16 Gyr with Galactic GC observations by Burstein et al. (1984; 17 clusters), Covino et al. (1995; 17 clusters), Trager et al. (1998; 18 clusters), Puzia et al. (2002; 12 clusters) and Beasley et al. (2004; 12 clusters). Note that some clusters were repeatedly observed, so more than one data point in the figure can refer to the same cluster. The metallicities are taken from Harris (1996, revision Feb. 2003). In Fig. 2 we show the time evolution of the model indices for all six metallicities, and confront them with Galactic GC observations (taken from the same references as in Fig. 1). The GC age determinations

<sup>1</sup> In view of the resolution of our spectral library, these values are not very accurate; however, since  $F_C$  is merely an additional weighting factor for the integration routine, this does not affect the final results.



**Fig. 1.** Indices  $H\beta$  (left) and  $Fe5335$  (right) versus metallicity for 5 different ages. Also shown are Galactic GC observations from various authors as indicated in the right-hand panel; GC metallicities are taken from Harris (1996, revision Feb. 2003). A typical measurement error is about 0.2 Å.



**Fig. 2.** Indices  $H\beta$  (left) and  $Fe5335$  (right) versus age for 6 different metallicities. Also shown are Galactic GC observations from various authors as indicated in the right-hand panel; GC age determinations are taken from Salaris & Weiss (2002).

are based on CMD fits and taken from Salaris & Weiss (2002)<sup>2</sup>. Over the range of Galactic GC ages and metallicities (i.e., ages older than about 8 Gyr and metallicities lower than solar in most cases), a sufficient agreement is observed between models and data in the sense that the data lie within the range of the model grid; we also checked this for other indices (not plotted).

However, the plots also demonstrate how difficult it would be to interpret the indices in terms of classical index-index plots. Actually, Fig. 2 seems to show apparent inconsistencies. So, some clusters in Fig. 2 have metallicities up to  $[Fe/H]=+0.4$

when compared with models for the age sensitive index  $H\beta$ , whereas, when compared with models for the metallicity sensitive  $Fe5335$  index, all clusters have metallicities lower than  $[Fe/H]=-0.4$ . We cannot decide at this point to what degree these inconsistencies are due to problems in the *models* or the calibrations they are based on or due to badly calibrated *observations*; however, our new Lick Index Analysis Tool nevertheless gives surprisingly robust age and, particularly, metallicity determinations for the same set of cluster observations (cf. Sect. 4.2).

<sup>2</sup> Note that they only cover a subsample of the observations shown in Fig. 1: 11 clusters of the Burstein et al. (1984) sample, 10 of the Covino et al. (1995) sample, 10 of the Trager et al. (1998) sample, only 3 of the Puzia et al. (2002) sample, and 6 clusters of the Beasley et al. (2004) sample.

### 3. Index sensitivities

It is well known that different indices have varying sensitivities to age and/or metallicity. To quantify this, Worthey (1994) in-

**Table 1.** Metallicity sensitivity parameters for different zero points. Low numbers indicate high age-sensitivity. Values given in brackets are not reliable (see text).

	WORTHY	GALEV 12 Gyr		GALEV 4 Gyr	
		Z = 0.02	0.0004	0.02	0.0004
CN <sub>1</sub>	1.9	1.5	1.1	1.1	0.1
CN <sub>2</sub>	2.1	1.5	0.2	1.2	0.1
Ca4227	1.5	1.1	(0.4)	1.0	0.1
G4300	1.0	1.0	0.2	0.8	0.1
Fe4383	1.9	1.9	0.3	1.3	0.2
Ca4455	2.0	1.7	0.7	1.5	0.3
Fe4531	1.9	1.7	0.6	1.4	0.2
Fe4668	4.9	(3.5)	(0.9)	2.4	0.9
H $\beta$	0.6	0.6	0.1	0.5	0.1
Fe5015	4.0	(2.3)	(1.3)	2.1	0.4
Mg <sub>1</sub>	1.8	1.7	(2.2)	1.4	2.0
Mg <sub>2</sub>	1.8	1.5	(1.8)	1.2	0.5
Mg <sub>b</sub>	1.7	1.4	(0.7)	1.0	0.3
Fe5270	2.3	2.0	(0.7)	1.6	0.3
Fe5335	2.8	2.7	(1.3)	2.0	0.4
Fe5406	2.5	(2.6)	(2.3)	1.8	0.6
Fe5709	6.5	(8.5)	(1.7)	2.6	(1.2)
Fe5782	5.1	(5.9)	(1.4)	2.5	(1.0)
Na D	2.1	1.9	(1.2)	1.9	0.6
TiO <sub>1</sub>	1.5	0.9	0.7	(1.4)	(5.5)
TiO <sub>2</sub>	2.5	1.3	0.9	(1.6)	(8.6)
H $\delta_A$	1.1	1.0	0.3	0.8	0.1
H $\gamma_A$	1.0	1.0	0.2	0.8	0.1
H $\delta_F$	0.9	0.9	0.1	0.7	0.1
H $\gamma_F$	0.8	0.8	0.2	0.7	0.1

roduced a “metallicity sensitivity parameter” that gives a hint on how sensitive a given index is with respect to changes in age and metallicity. This parameter is defined as the ratio of the percentage change in Z to the percentage change in age (so influences of possible age-metallicity degeneracies are implicitly included), with large numbers indicating greater metallicity sensitivity:

$$S = \left( \frac{\Delta I_Z}{\Delta Z/Z} \right) / \left( \frac{\Delta I_{age}}{\Delta age/age} \right) \quad (6)$$

Using his SSP models, Worthey (1994) chose a 12 Gyr solar metallicity (Z=0.017) model as the zero point for the sensitivity parameters, the  $\Delta$ ’s referring to “neighbouring” models, in this case models with age = 8/17 Gyr (i.e.,  $\Delta age = 4/5$  Gyr) and Z = 0.01/0.03 (i.e.,  $\Delta Z = 0.007/0.013$ )<sup>3</sup>; the main numerator/denominator in Eq. 6 is averaged using both  $\Delta$ ’s before computing the fraction.

In Table 1, we reprint the metallicity sensitivity parameters given by Worthey (1994) and Worthey & Ottaviani (1997), and confront them with parameters computed using our own models. Extending Worthey’s approach, we compute parameters for four different combinations of zero points, using high (Z = 0.02) and low (Z = 0.0004) metallicities, and high (12 Gyr) and intermediate (4 Gyr) ages.

<sup>3</sup> Ideally, S should be relatively independent of the exact values of the  $\Delta Z$  and  $\Delta age$  chosen, as long as they are not too large.

**Table 2.** “Neighbouring models” in terms of metallicity for Z=0.02 and Z=0.0004 (*top*) and in terms of age for the zero points 12 Gyr and 4 Gyr (*bottom*) used to compute the metallicity sensitivities given in Table 1. Brackets give the corresponding  $\Delta Z$  and  $\Delta age$ , respectively.

Z=0.02		Z=0.0004	
0.008(0.012)	0.05(0.03)	0.0001(0.0003)	0.004(0.0036)

12 Gyr		4 Gyr	
11.0(1.0)	13.2(1.2)	3.2(0.8)	5.0(1.0)
10.5(1.5)	13.8(1.8)	2.5(1.5)	6.3(2.3)

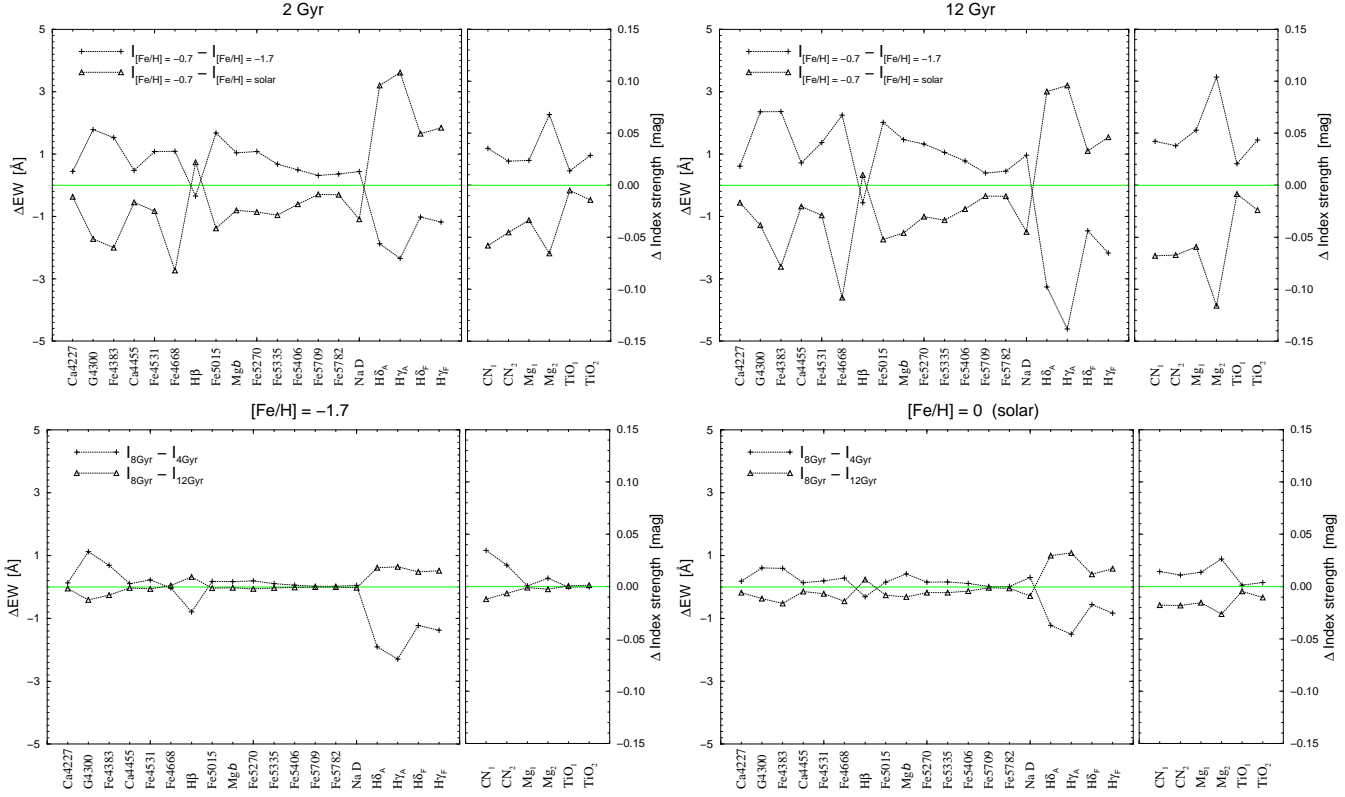
Worthey’s parameters are relatively well reproduced by models with a similar combination of zero points, i.e. for old (12 Gyr) and solar metallicity SSPs; however, S is not totally independent of the  $\Delta Z$  and  $\Delta age$  chosen; it can be very sensitive to the exact evolution of the model index, mainly in age-metallicity space regions where the slope of the index does not evolve very smoothly (for example, a very high value of the parameter can also mean that, due to a small “bump” in the time evolution of the index,  $\Delta I_{age}$  is near zero; in this case, S is worthless). The two zero points for both age and metallicity and their “neighbouring models” we use for the computations are given in Table 2; for both age zero points, we chose two sets of neighbouring models and averaged the final parameters. In order to check the reliability of our results, we also computed parameters for values of age and  $\Delta age$  not given in the table. If the results for different  $\Delta age$ ’s (or slightly different zero points) differ strongly, we classify the parameter as uncertain (indicated by brackets in Table 1).

The “ranking” of indices in terms of sensitivity is, with some exceptions, unaffected by changes in the zero points.

However, contrary to our expectations, at solar metallicity the age-sensitivity of Lick indices is only slightly higher for an intermediate age model compared to the 12 Gyr model; for low-metallicity SSPs, the effect is more pronounced. Most important, we find the result that for models at low metallicity, indices are generally much more age sensitive than for models at high metallicity, especially for age sensitive indices like G4300 or Balmer line indices. This means, indices of old, low metallicity GCs can be more sensitive to age than indices of GCs with high metallicity and intermediate age. This is of special interest for any analysis of GC systems involving intermediate age GCs (e.g., in merger remnants), since secondary GC populations with intermediate ages are generally expected to have higher metallicities than “normal” old and metal-poor populations.

Given the limited accuracy of any index measurement, in practice the usefulness of an index to determine age or metallicity does not only depend on the *relative* change in index strength for changing Z or age as it is given by S but also on the *absolute* change in index strength.

Therefore, in Figure 3 we show the absolute differences of index strengths for old (12 Gyr) and young (2 Gyr) SSPs for changing metallicity, and for metal-rich ([Fe/H]=0) and



**Fig. 3.** Absolute differences of index strengths for old and young SSP models for changing metallicity (*top*), and for metall-rich and metall-poor SSP models for changing age (*bottom*). The dotted lines are just for presentation.

metall-poor ( $[\text{Fe}/\text{H}] = -1.7$ ) SSPs for changing age, respectively. Generally, the absolute differences between 8 and 4 Gyr old SSPs are larger than the differences between 8 and 12 Gyr old SSPs at fixed metallicity, as expected (Fig. 3, lower panels). However, this effect is much stronger at low than at high metallicity, which confirms what we get from the  $S$  parameter. The absolute differences between models with different metallicity (top panels in Fig. 3) are slightly larger for old than for young SSPs. Interestingly, the plots show that indices known to be sensitive to age can also be highly variable for differing metallicities; especially the broad Balmer indices  $H\delta_A$  and  $H\gamma_A$  change strongly with metallicity. Most important, however, the plot shows that in practice, moderately metal-sensitive indices like  $\text{Mgb}$  can be much more useful for metallicity determinations than indices like  $\text{Fe5709}$  or  $\text{Fe5782}$ , though the latter are, according to the  $S$  parameter, much more metal-sensitive.

In order to determine ages and metallicities of GCs, indices should be chosen not only according to known sensitivities as given by  $S$ , but also according to the achievable measurement accuracy and, if possible, according to the expected age and metallicity range of the sources.

#### 4. The Lick Index Analysis Tool

Since in the original models (cf. Sect. 2.1), the steps in metallicity are large, in a first step we linearly interpolate in  $[\text{Fe}/\text{H}]$  between the 6 metallicities before we analyse any data with our new tool. This is done in steps of  $[\text{Fe}/\text{H}] = 0.1$  dex, so the final

input grid for the analysis algorithm consists of sets of all 25 Lick indices each for 28 metallicities ( $-2.3 \leq [\text{Fe}/\text{H}] \leq +0.4$ ) and 4000 ages ( $4 \text{ Myr} \leq \text{age} \leq 16 \text{ Gyr}$ ). Although this approach is only an approximation, the results shown in Sect. 4.2 prove it to be sufficiently accurate.

##### 4.1. The $\chi^2$ - approach

The algorithm is based on the SED Analysis Tool presented by Anders et al. (2004); the reader is referred to this paper for additional information about the algorithm, as well as for extensive tests using broad-band colors instead of indices.

All observed cluster indices at once – or an arbitrary subsample of them – are compared with the models by assigning a probability  $p(n)$  to each model grid point (i.e., to each set of 25 indices defined by 1 age and 1 metallicity):

$$p(n) \propto (-\chi^2), \quad (7)$$

where

$$\chi^2 = \sum_{i=1}^{25} \frac{(I_{\text{obs}} - I_{\text{model}})^2}{\sigma_{\text{obs}}^2 + \sigma_{\text{model}}^2} \quad (8)$$

with  $I_{\text{obs}}$  and  $I_{\text{model}}$  being the observed and the model indices, respectively, and  $\sigma_{\text{obs}}$  and  $\sigma_{\text{model}}$  being the respective uncertainties. Indices measured in magnitudes are transformed into Ångström before calculation. After normalization ( $\sum p(n) = 1$ ), the grid point with the highest probability is

**Table 3.** Observations by Burstein et al. (1984, B84), Covino et al. (1995, C95), Trager et al. (1998, T98;  $H\delta$ ,  $H\gamma$  are taken from Kuntschner et al. 2002, see text), Beasley et al. (2004, B04) used to perform the tests of Section 4.2. ‘\*’: index observed, ‘o’: only a subsample of clusters is observed in this index.

	B84	C95	T98	B04
CN <sub>1</sub>	*	o	*	o
CN <sub>2</sub>			*	o
Ca4227			*	o
G4300	*	o	*	o
Fe4383			o	o
Ca4455			*	o
Fe4531			*	o
Fe4668			*	*
$H\beta$	*	*	*	*
Fe5015			*	*
Mg <sub>1</sub>	*	*	*	*
Mg <sub>2</sub>	*	*	*	*
Mgb	*	*	*	*
Fe5270	*	*	*	*
Fe5335	*	*	*	*
Fe5406			*	*
Fe5709			*	*
Fe5782			o	*
Na D	*	o	*	*
TiO <sub>1</sub>	*		*	*
TiO <sub>2</sub>	o		o	
$H\delta_A$			*	o
$H\gamma_A$			o	o
$H\delta_F$			*	o
$H\gamma_F$			*	o

assumed to be the *best model*, i.e. it gives the “best age” and the “best metallicity” for the observed cluster.

The uncertainties of the best model in terms of  $\pm 1\sigma$  *confidence intervals* are computed by rearranging the model grid points by order of decreasing probabilities, and summing up their probabilities until  $\sum p(n) = 0.68$  is reached; the  $1\sigma$  uncertainties in age and metallicity are computed from the age and metallicity differences, respectively, of the  $n(p_{0.68})$ - and the  $n(p_{max})$ -model. Note that the determination does not take into account the possible existence of several solution “islands” for one cluster; thus the confidence intervals are in fact upper limits.

#### 4.2. Examples and tests I: Galactic GCs

We have tested our Lick Index Analysis Tool using a large set of Galactic GCs for which index measurements (taken from Burstein et al. 1984, Covino et al. 1995, Trager et al. 1998<sup>4</sup>, and Beasley et al. 2004) as well as age and metallicity determinations from CMD analyses (taken from Salaris &

<sup>4</sup> In this dataset,  $H\delta_A$ ,  $H\gamma_A$ ,  $H\delta_F$  and  $H\gamma_F$  are taken from Kuntschner et al. 2002 who reanalysed Trager et al.’s spectra; in the following, ‘Trager et al. 1998’ always is meant to include this additional data.

**Table 4.** Mean ages and standard deviations of cluster ages determined using the Lick Index Analysis Tool and CMD analysis (Salaris & Weiss 2002), respectively, as shown in Figs. 4 and 5. Note that the values are computed without cluster NGC 6121.

	all indices		age-sensitive indices	
	<age>	$\pm$	<age>	$\pm$
Lick–Analysis	13.49	1.80	13.58	1.22
CMD–Analysis	11.54	1.08	11.57	1.04

Weiss 2002) are available.

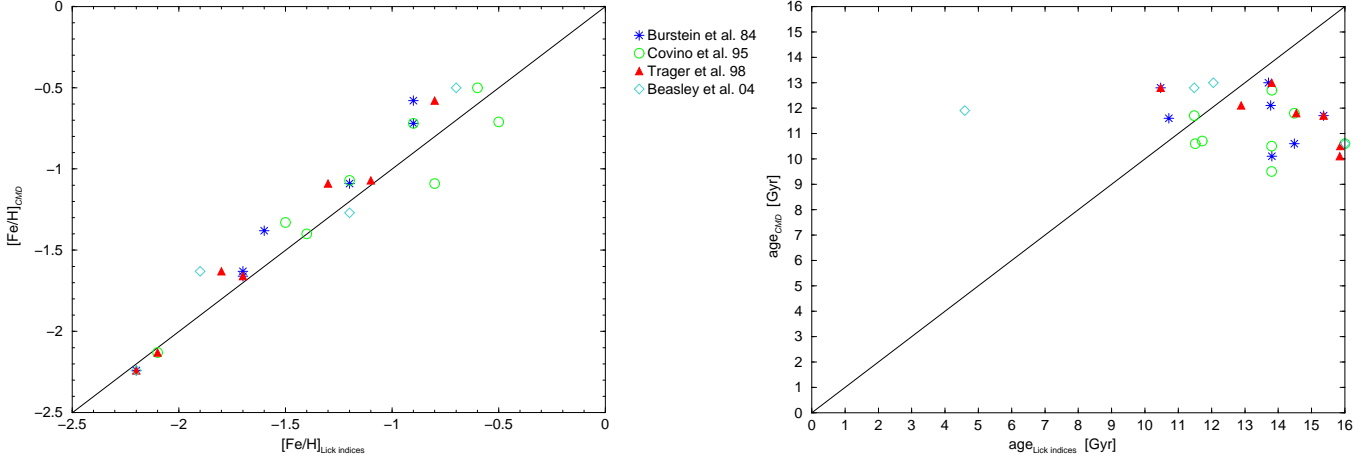
Fig. 4 compares ages and metallicities from both methods. Here, we use the complete set of measured indices available (cf. Table 3) as input for our Analysis Tool; for comparison, Fig. 5 shows our results using two subsets of indices: the age-sensitive indices Ca4227, G4300,  $H\beta$  and TiO<sub>1</sub> on the left panel, and metal-sensitive indices Mg<sub>1</sub>, NaD, [MgFe] plus the age-sensitive index  $H\beta$  on the right panel<sup>5</sup>. In all plots, only results with confidence intervals of  $\sigma(\text{age}) \leq 5$  Gyr are plotted<sup>6</sup>.

The agreement between [Fe/H] obtained from our Lick Index Analysis Tool and the corresponding values from CMD analyses is very good, with  $\Delta[\text{Fe}/\text{H}] \leq 0.3$  dex when using all available indices, and  $\Delta[\text{Fe}/\text{H}] \leq 0.2$  dex when using mainly metal-sensitive indices. With one exception, the age determinations are relatively homogeneous, though the mean age obtained from index analyses is about 2 Gyrs too high compared to the results from CMD analyses. Table 4 gives the mean ages and standard deviations of clusters determined using the Lick Index Analysis Tool and from CMD analyses, respectively. It shows that, using all available indices, not only the mean ages but also the age spreads are too high; most likely, this is due to varying Horizontal Branch (HB) morphologies (see below). However, if only age-sensitive indices are used, the age spread is of the same magnitude than that obtained by CMD analyses.

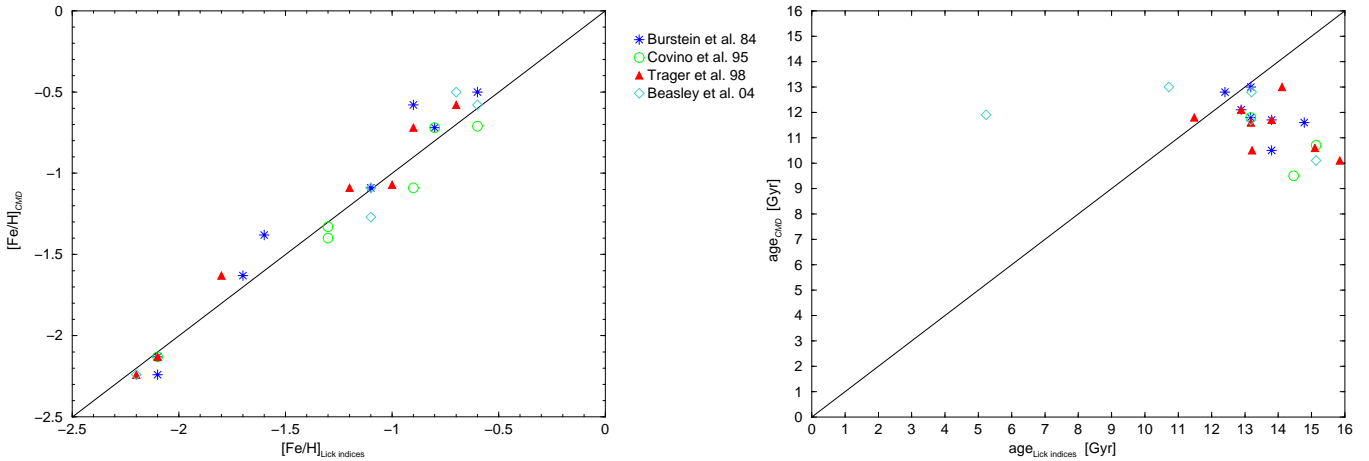
As an example, Figure 6 shows the “best-fitting” model for the Galactic GC M3 (NGC 5272) together with the index measurements of Trager et al. (1998) used for the analysis. The best model has an age of  $12.88^{(-1.99)}_{(+1.75)}$  Gyr and a metallicity of  $[\text{Fe}/\text{H}] = -1.7(\pm 0)$  dex; compared with an age of  $12.1(\pm 0.7)$  Gyr and  $[\text{Fe}/\text{H}] = -1.66$  dex given by CMD analysis, this is a very good solution. Note that the dotted line connecting the model indices is just for presentation. We also give the  $\pm 1\sigma$  confidence intervals of our best model in terms of index values for SSPs with age  $12.89 - 1.99 = 10.90$  Gyr and  $12.89 + 1.74 = 14.63$  Gyr, respectively, and metallicity  $[\text{Fe}/\text{H}] = -1.7$ .

<sup>5</sup> [MgFe] is a combination of metal-sensitive indices which is known to be widely unaffected by non-solar abundance ratios (see, e.g., Thomas et al. 2003). It is defined as  $[\text{MgFe}] := \sqrt{\langle \text{Fe} \rangle \times \text{Mgb}}$ , with  $\langle \text{Fe} \rangle := (\text{Fe5270} + \text{Fe5335}) / 2$ .

<sup>6</sup> In most cases, very large  $1\sigma$  uncertainties are due to the presence of two “solution islands” (e.g., solution 1: low or intermediate age, solution 2: high age) which are both within their  $1\sigma$  ranges. Since we do not want to use any a priori information about the clusters, we cannot decide between the two solutions and therefore rather omit them completely.



**Fig. 4.** Galactic GC observations: Metallicities (*left*) and ages (*right*) determined using our Lick Index Analysis Tool (x-axis, using all measured indices available) vs. metallicities and ages determined by CMD analyses (y-axis, taken from Salaris & Weiss 2002). Note that only results with *model* uncertainties of  $\sigma(\text{age}) \leq 5$  Gyr are plotted.



**Fig. 5.** Same as Figure 4, but using metallicity-sensitive indices  $\text{Mg}_1$ ,  $\text{NaD}$ ,  $[\text{MgFe}]$  plus age-sensitive index  $\text{H}\beta$  as input only (*left*), and using age-sensitive indices  $\text{Ca}4227$ ,  $\text{G}4300$ ,  $\text{H}\beta$ , and  $\text{TiO}_1$  as input only (*right*).

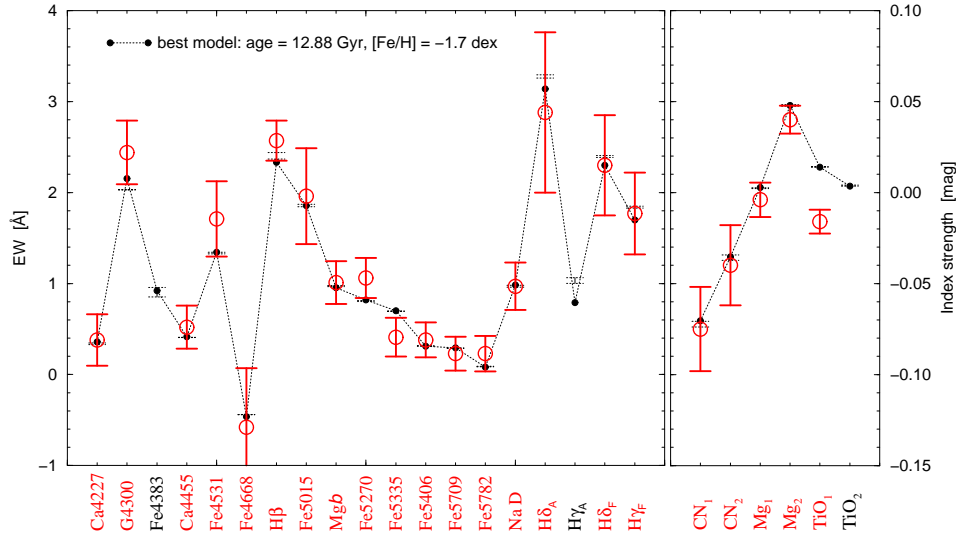
As seen in Figs. 4 and 5, most of Galactic GCs are very well recovered in their metallicities by our Lick Index Analysis Tool, in particular when the analysis is concentrated on the metal-sensitive indices  $\text{Mg}_1$ ,  $\text{NaD}$ ,  $[\text{MgFe}]$  and the age-sensitive index  $\text{H}\beta$ . The origin of the  $\sim 2$  Gyr systematic difference between index-determined and CMD-based ages, as well as of the wider age spread we find is, most likely, due to the HB morphologies of the clusters: The Padova isochrones we use for the analyses have very *red* HBs over most of the parameter space (they have blue HBs only for metallicities  $[\text{Fe}/\text{H}] \leq -1.7$  and ages higher than about 12 Gyr); therefore, the age of an observed cluster with *blue* HB can possibly be underestimated by several Gyrs. Proctor et al. (2004), who use a similar technique than that applied here, also find ages too high compared to values from CMD analyses; depending on the applied SSP models, they find mean ages of  $13.1(\pm 2.3)$ ,  $12.2(\pm 3.3)$  and  $12.7(\pm 1.9)$  Gyr, respectively (cf. with Tab. 4). We plan to anal-

yse the influence of HB morphology on Lick index based age determinations in a separate paper.

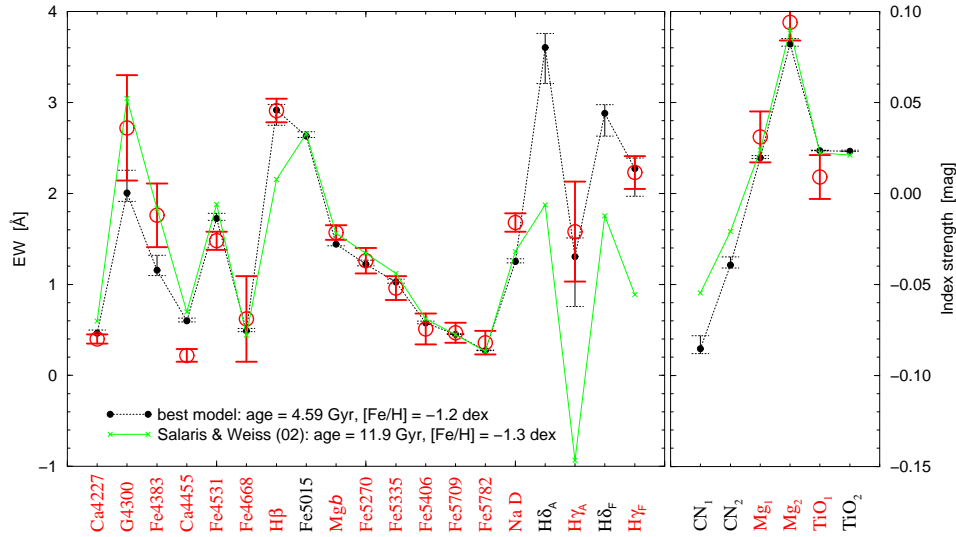
Interestingly, and despite the fact that the Lick index measurements used here are of very different age and quality, the results are of comparable quality for each data set. E.g., the indices taken from Trager et al. (1998) are measured using the same original Lick-spectra than the Burstein et al. (1984) data set; however, the spectra were recalibrated, and more indices were measured. Nonetheless, the results from both data sets are comparable.

Though most results are acceptable, one cluster of our set is seriously misdetermined in terms of age: For the Galactic GC M4 (NGC 6121) the Lick Index Analysis Tool gives an age of only  $\sim 5$  Gyr (with a  $1\sigma$  uncertainty of less than 1 Gyr) both using all and only age-sensitive indices; CMD analysis gives more than twice the age. Since the cluster does *not* have a very blue HB (Harris 1996 gives a HB ratio of nearly zero), we do





**Fig. 6.** Lick index measurements of the Galactic GC M3 (NGC 5272) by Trager et al. (1998) with observational errors (open circles), and “best model” indices with the  $\pm 1\sigma$  confidence intervals (black dots). The best model has an age of  $12.88^{(-1.99)}_{(+1.75)}$  Gyr and  $[\text{Fe}/\text{H}] = -1.7(\pm 0)$  dex; Salaris & Weiss (2002) give age =  $12.1(\pm 0.7)$  Gyr and  $[\text{Fe}/\text{H}] = -1.66$  dex.



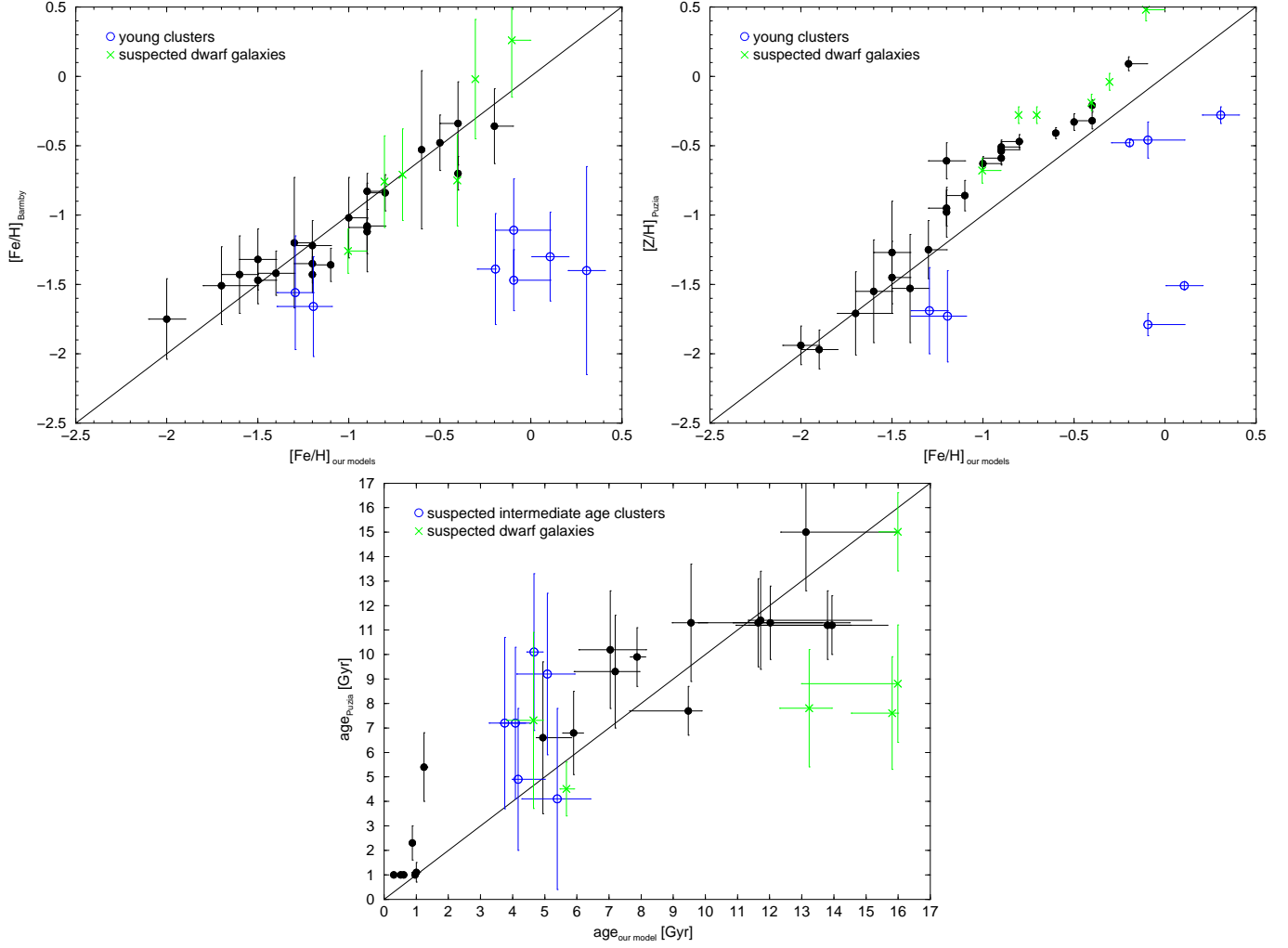
**Fig. 7.** Lick index measurements of the Galactic GC M4 (NGC 6121) by Beasley et al. (2004) with observational errors (open circles), and “best model” indices with the  $\pm 1\sigma$  confidence intervals (black dots). The best model has an age of only  $4.59^{(-0.31)}_{(+0.80)}$  Gyr and  $[\text{Fe}/\text{H}] = -1.2^{(-0.1)}_{(+0.0)}$  dex; Salaris & Weiss (2002) give age =  $11.9(\pm 1.1)$  Gyr and  $[\text{Fe}/\text{H}] = -1.27$  dex. Additionally, we plot model indices for the Salaris & Weiss (2002) solution, i.e., a  $11.9$  Gyr /  $[\text{Fe}/\text{H}] = -1.3$  dex SSP model (small crosses).

not have a reasonable explanation for this. However, anomalies have been found for this cluster, and some properties are still discussed in the literature (see, e.g., Richer et al. 2004 and references therein). Figure 7 shows models for this “misdetermined” cluster: Together with the index measurements taken from Beasley et al. (2004), we show the index values for our best model (i.e., indices for a SSP with age =  $4.59^{(-0.31)}_{(+0.80)}$  Gyr and  $[\text{Fe}/\text{H}] = -1.2^{(-0.1)}_{(+0.0)}$  dex) as well as for a model SSP using the Salaris & Weiss (2002) solution (age =  $11.9(\pm 1.1)$  Gyr and  $[\text{Fe}/\text{H}] = -1.3$  dex). The indices which differ most between the two models (and for which measurements are available) are

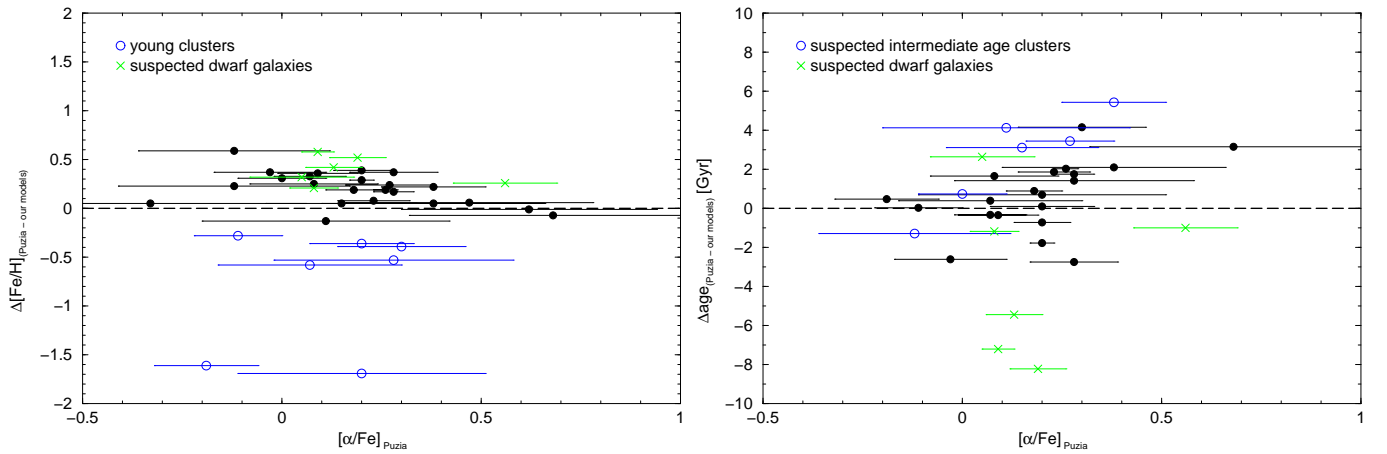
G4300, Fe4383, and the Balmer line indices  $\text{H}\beta$  and  $\text{H}\gamma$ ; remarkably, the Balmer lines seem to be completely responsible for the misdetermination.

#### 4.3. Examples and tests II: M31 GCs and non-solar abundance ratios

Unfortunately, for Andromeda galaxy (M31) GCs it is not possible to obtain high quality color magnitude diagrams; reliable determinations of age and metallicity which could be used as “default values” for comparisons are not available. Therefore,



**Fig. 8.** M31 GCs: Metallicities and ages for the Beasley et al. (2004) GC sample determined using our Lick Index Analysis Tool (x-axis, using all measured indices available) vs. metallicity determinations taken from Barmby et al. (2000) (*top left*) and Puzia et al. (2005) (*top right*), and vs. age determinations taken from Puzia et al. (2005) (*bottom panel*). The classification as “young cluster” and “suspected dwarf galaxy” is taken from Beasley et al. (2004). See electronic edition for color versions of these plots.



**Fig. 9.** M31 GCs: Absolute differences between parameters derived using our Lick Index Analysis Tool and results from Puzia et al. (2005), against  $[\alpha/\text{Fe}]$  taken from Puzia et al. (2005). *Left:*  $([\text{Z}/\text{H}]_{\text{Puzia}} - [\text{Fe}/\text{H}]_{\text{our models}})$ ; *Right:*  $(\text{age}_{\text{Puzia}} - \text{age}_{\text{our models}})$ . The classification as “intermediate-age cluster” and “suspected dwarf galaxy” is taken from Beasley et al. (2004).

for M31 GCs we can only compare our Lick index based determinations with results taken from the literature which are based on spectral indices themselves.

For our analyses, we use the Lick index measurements of M31 GCs presented by Beasley et al. (2004); while not presenting own age or metallicity determinations for individual clusters, they distinguish four classes for their sample of cluster candidates: Young, intermediate age, and “normal” old GCs. Additionally, some sources are suspected to be foreground galaxies. Beasley et al. have measured all available Lick indices with the exception of  $\text{TiO}_2$ .

In Figure 8, we compare our metallicity determinations using the Lick Index Analysis Tool with results presented by Barmby et al. (2000) (top left panel) and Puzia et al. (2005) (top right panel). While Barmby et al. use calibrations given by Brodie & Huchra (1990) for their spectroscopic metallicity determinations, using own measurements of absorption line indices, Puzia et al. (2005) use a  $\chi^2$  approach using Lick index models from Thomas et al. (2003, 2004), which do account for non-solar abundance ratios. Puzia et al. use the same database as we do (i.e., the Lick index measurements published by Beasley et al. 2004). They give not  $[\text{Fe}/\text{H}]$  but total metallicities  $[\text{Z}/\text{H}]$ ; however, according to Thomas et al. (2003),  $[\text{Fe}/\text{H}]$  in the ZW84 scale is in excellent agreement with  $[\text{Z}/\text{H}]$ . Hence, our results, given in  $[\text{Fe}/\text{H}]$ , are perfectly comparable to Puzia et al.’s results, and well appropriate to test for the influence of non-solar abundance ratios on our results. For both the Barmby et al. (2000) and Puzia et al. (2005) metallicity determinations, we find good agreement with our results. Only for clusters which are classified as young (i.e., age  $\leq 1\text{-}2$  Gyr) we find relatively large differences in  $[\text{Fe}/\text{H}]$ ; however, this reflects our expectations, since the models are calibrated using intermediate-age or old Galactic stars mainly.

In the bottom panel of Fig. 8, we confront our results with ages determined by Puzia et al. (2005). Again, the results are in surprisingly good agreement, if sources suspected to be foreground dwarf galaxies are not considered. For the set of intermediate-age clusters identified by Beasley et al. (2004), our results perfectly reflect this classification.

Compared with the classification of Beasley et al. (2004), the largest disagreements in both age and metallicity occur for young clusters, and for suspected dwarf galaxies; it is no surprise that models computed to fit GCs are not appropriate to be applied to galaxies (and, therefore, different methods lead to different results).

Since Puzia et al. (2005) also determine  $\alpha$ -enhancements for the GC sample, we can check for possible systematic offsets of our determinations compared to theirs due to non-solar abundance ratios.

In Figure 9, absolute differences between metallicities (left panel) and ages (right panel) derived using our models and from Puzia et al. (2005) are plotted against  $[\alpha/\text{Fe}]$ . Relatively surprising, no general trend for the differences in both age and metallicity determinations with  $\alpha$ -enhancement can be observed, if the large error bars of the  $[\alpha/\text{Fe}]$  determinations are taken into account. Hence, the slight offset between metallicities

determined by Puzia et al. and by us (cf. Fig. 8, top right panel) for  $[\text{Fe}/\text{H}]$  larger than  $\sim -1.2$  dex seems *not* to be due to the use of solar-scaled against  $\alpha$ -enhanced models.

## 5. Summary and outlook

To cope with the observational progress that makes star cluster & globular cluster spectra accessible in a large variety of external galaxies, we have computed a large grid of evolutionary synthesis models for simple stellar populations, including 25 Lick/IDS indices using the empirical calibrations of Worthey et al. (1994) and Worthey & Ottaviani (1997). Comparison of the SSP models with Galactic GC observations shows good agreement between models and data.

We find that the well-known and widely used age-sensitive indices  $\text{H}\delta_A$  and  $\text{H}\gamma_A$  also show a strong metallicity dependence. The “metallicity sensitivity parameter”  $S$  introduced by Worthey (1994) for old stellar populations with solar metallicity is well reproduced by our models. Our models allow to extend this concept towards younger ages and non-solar metallicities. We find the sensitivity of different indices with respect to age and metallicity to depend itself on age and metallicity. E.g., all indices are generally more age sensitive at low than at high metallicity. Another important issue is the *absolute* difference in index strength for varying age or metallicity: Due to the limited accuracy of any index measurement, these absolute differences in practice can be of higher importance than the sensitivity given by  $S$ .

We present a new advanced tool for the interpretation of absorption line indices, the Lick Index Analysis Tool *LINO*. Following a  $\chi^2$ -approach, this tool determines age and metallicity including their respective  $\pm 1\sigma$  uncertainties using all, or any subset, of measured indices. Testing our tool against index measurements from various authors for Galactic GCs, which have reliable age and metallicity determinations from CMD analyses in the literature, shows very good agreement: Metallicities of GCs are recovered to  $\pm 0.2$  dex using 6 appropriate indices only ( $\text{Mg}_1$ ,  $\text{Mgb}$ ,  $\text{Fe}5270$ ,  $\text{Fe}5335$ ,  $\text{NaD}$ ,  $\text{H}\beta$ ). Age determinations from Lick indices consistently yield ages  $\sim 2$  Gyr higher than those obtained from CMDs. The origin of this discrepancy is not yet understood. Index measurements for M31 clusters are analysed and compared to results from the literature; a good agreement between our results and age and metallicity determinations from the literature is found. We show that the drawback of not having non-solar abundance ratio models do not seriously affect our results.

We will apply *LINO* to the interpretation of intermediate-age and old GC populations in external galaxies, complementing our SED Analysis Tool for the interpretation of broad-band spectral energy distributions.

All models are accessible from our website, <http://www.astro.physik.uni-goettingen.de/~galev/>.

*Acknowledgements.* TL is partially supported by DFG grant Fr 916/11-1-2-3.

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